

# Exploiting the Magnetic Origin of Solar Activity in Forecasting Thermospheric Density Variations

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**Abstract** A detailed understanding of solar irradiance and its variability at extreme ultraviolet (EUV) wavelengths is required to model thermospheric density and to specify and forecast satellite drag. Current operational models rely on forecasts of proxies for solar activity based on autoregression. The forecasts from these models generally degrade to climatology after only a few days. Solar magnetic fields are ultimately responsible for variations in the EUV irradiance. The evolution of solar magnetic fields is well understood and results from a combination of solar rotation, diffusion, meridional flow, and magnetic flux emergence. In this presentation we review the current state of autoregressive proxy models and compare their forecast skill against new activity models based on magnetic flux transport.

## 1. Introduction

The influence of the Sun's extreme ultraviolet radiation (EUV, 50–1200 Å) on the Earth's upper atmosphere is perhaps the most dramatic of all the Sun-Earth connections. The solar irradiance at these wavelengths is directly coupled to solar activity and is strongly modulated by the 11-year sunspot cycle. Solar EUV radiation is deposited in the Earth's atmosphere at heights above about 100 km where the densities are low. The coupling of the highly variable solar EUV emission to the Earth's tenuous upper atmosphere leads to dramatic variations in the thermodynamics, chemical composition, and ionization of the Earth's upper atmosphere over the solar cycle. The maximum temperature of the thermosphere, for example, typically rises by about 500 K (750 to 1250 K) from solar minimum to solar maximum. Both the neutral and electron density in this region can change by more than an order of magnitude during a typical solar cycle [for reviews see 6, 9].

The close connection between solar variability at these wavelengths and thermospheric density means that accurately specifying and forecasting satellite drag requires a detailed understanding of the solar EUV irradiance. Unfortunately, establishing and monitoring the absolute radiometric calibration of EUV instruments is difficult and much of our quantitative understanding of the solar EUV irradiance and its variability remains uncertain. For example, it is not clear if the unusually low values for various irradiance measurements taken during the recent solar minimum reflect real changes in the solar atmosphere or are simply instrumental problems. The Solar EUV Monitor (SEM) on SOHO recorded a 15% reduction in the irradiance near 304 Å between the 1996 and 2008 minima [18], but the behavior of F10.7 and the Mg II core-to-wing ratio, proxies for solar activity, suggests that the irradiance during the two solar minima should differ by only a few percent [2, 14]. In the absence of continuous and well calibrated solar spectral irradiance measurements, models of thermospheric density rely on proxies for solar activity.

In this paper we explore the use of solar magnetic field measurements as both the basis of solar EUV irradiance variability models as well as for use in forecasting irradiance variability. Spatially resolved measurements of solar radiance are well correlated with the magnitude of the underlying magnetic field. This is the so-called “flux-luminosity” relationship [e.g., 12, 3], which holds in conditions ranging from the quiet Sun to the most active stars [11]. This concept has already been applied to models of the total solar irradiance [e.g., the SATIRE-S model 1]) and is likely to form the basis of future irradiance models at other wavelengths. Furthermore, the evolution of solar magnetic fields is generally well understood and results from a combination of solar rotation, diffusion, meridional flow, and magnetic flux emergence [for a review

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see 10]. Thus it should be possible to make accurate forecasts of solar magnetic fields and proxies for solar activity related to them [see 4, for the application of this idea to forecasting F10.7 ]. Current forecasts of solar activity are often based on polynomial extrapolations (or autoregression) and their performance degrades quickly. Our goal is to make quantitative comparisons between models based on autoregression and models based on magnetic flux transport.

## 2. Autoregressive Forecast Models

The 27 day rotational period of the sun leads to a quasi-periodic modulation in the solar irradiance at all wavelengths. Current forecasts of proxies for solar activity often use autoregression to predict the amplitude of this quasi-periodic signature [e.g. 16, 8]. This approach makes use of a least-squares polynomial fit of “past” data to “future” data. That is, the observations are fit to an equation of the form  $f_{+n} = c + a_0 f_0 + a_{-1} f_{-1} + a_{-2} f_{-2} + \dots$ , where  $f_{+n}$  is the value of the proxy  $n$  days in the future from a given date and the right hand side of the equation represents the values of the proxy on the date of interest and before. Here we address several basic questions, such as what is the optimal amount of past data to include? How does the forecast error increase with the length of the forecast? and How does autoregression compare with simple reference models?

For this analysis we use the daily F10.7 observations collected over the past 6 solar cycles (January 1, 1954 to June 30, 2014). Note that of these 22,066 observations we have generated spline interpolated values for 168 (0.8%) missing observations. Most of these missing data points are from the 1950’s. This period is included because it represents the highest levels of solar activity observed over the past several hundred years. This data was randomly divided into a training set comprising 60% of the observations, and a validation set, which contains the remaining 40% of the data. The training data is used to determine the coefficients for the polynomial fit while the validation data is used to estimate the forecast error and skill.

We considered forecasts of 1 to 45 days using ensembles of between 1 and 108 past data values. For each model we computed the root-mean-square-error (RMSE) from the validation set as a function of forecast day and model complexity. The lowest forecast error is generally achieved for models with the most coefficients, although we found that there is little marginal benefit to considering more than about 30 previous observations. For the remainder of this paper we present results from models that use 108 previous data points. The RMSE as a function of forecast day for this case is shown in Fig. 1. Here we see that the RMSE rises rapidly for about 10 days and then asymptotes.

To evaluate the skill of the autoregressive forecast we compared it with two very simple reference forecast models: “persistence” — predicting that the last measured value will be observed for the next 45 days, and “climatology” — predicting that the average value from the previous 81 days will be observed over the next 45 days. We compute the model skill using

$$\text{Skill} = 1 - \frac{\text{MSE}(\text{model})}{\text{MSE}(\text{reference})}. \quad (1)$$

This comparisons shows the expected result, with the persistence reference model performing similarly to autoregression for very short periods, and autoregression converging to climatology over long periods. The peak skill in the autoregression model (0.4) is achieved at 6 days.

Note that the forecast and skill reported here are for the entire data set. These metrics vary with the solar cycle. At high levels of solar activity RMSE increases, but skill remains relatively high. At low levels of solar activity RMSE decreases, but since there is very little modulation in the irradiance and the simple reference models perform well, the skill of the forecast also falls.

We have repeated this exercise for the S10 proxy, which is simply a scaled version of the broad band irradiance measured by the SEM on SOHO [15]. These results, which are also shown in Fig. 1, are similar to those obtained for f10.7, except that the RMSE for the autoregressive forecast are about 50% smaller.

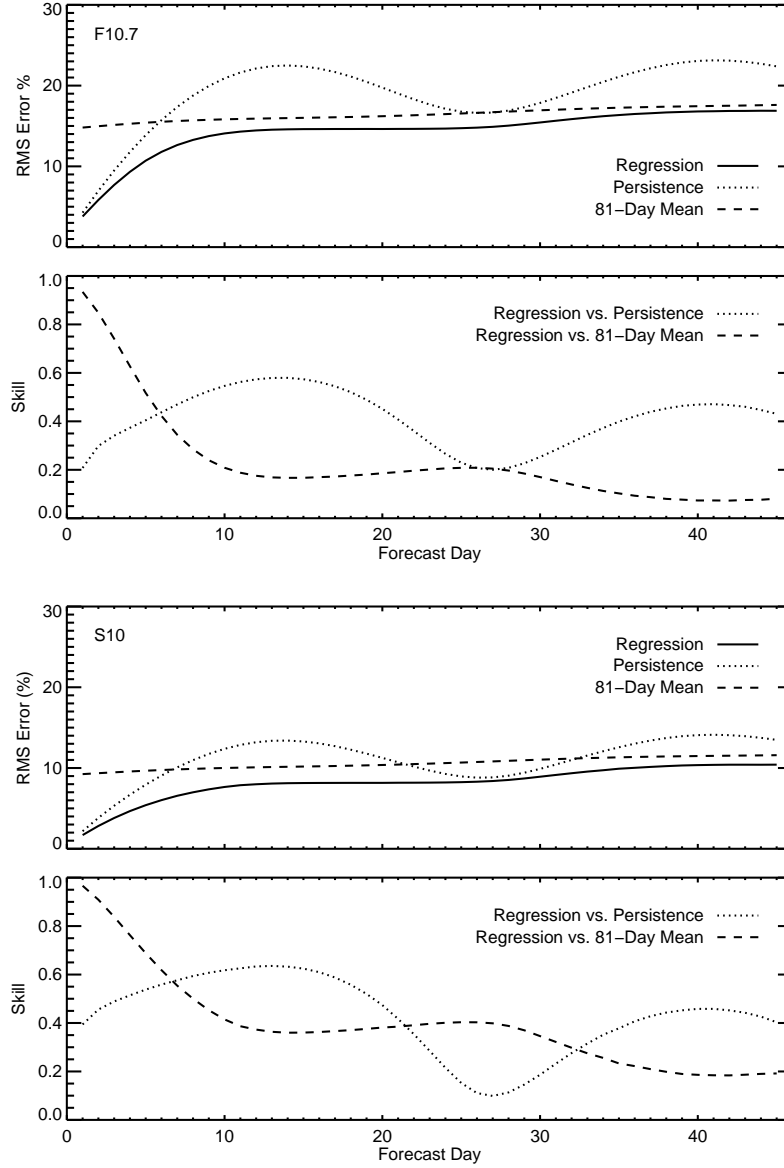


Figure 1: RMSE and forecast skill as a function of forecast day for autoregressive models of two proxies for solar activity. The F10.7 cm radio flux is shown in the top panels and S10 is shown in the bottom panels. The autoregressive models are compared against simple models of persistence and the 81-day running mean. For both proxies forecast skill peaks at about 6 days.

As we will demonstrate in the next section, this out-performance is likely due to the fact that S10 is more closely tied to emission from the lowest layers of the solar atmosphere.

### 3. Using the Solar Magnetic Field as a Proxy for Solar Activity

Several recent irradiance models of the total solar irradiance have used spatially resolved magnetic field measurements as an input [e.g., SATIRE-S 1, 5]. The basic assumption in these models is that only changes in the surface magnetic fields impact the solar irradiance on time scales greater than one day [e.g., 7]. Here

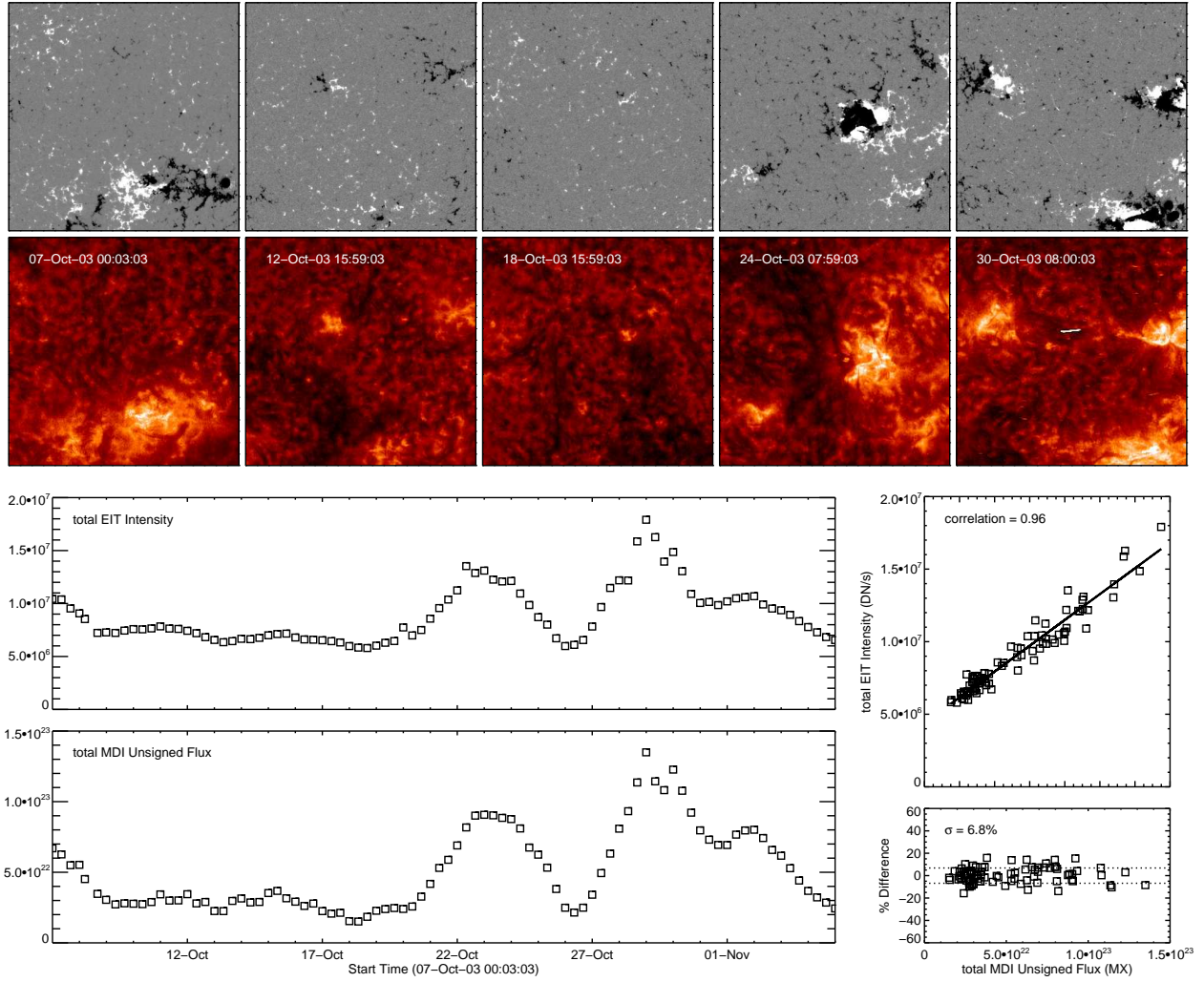


Figure 2: Observations of radiance and magnetic field at disk center. These observations illustrate the strong correlation between the integrated radiance and the total unsigned magnetic flux. The top panels show examples of the radial magnetic field measured with the Michelson Doppler Imager (MDI) and the 304 Å radiance from the EUV Imaging Telescope (EIT) measured at disk center. The bottom left panels show the corresponding time series. The bottom right panels show a linear regression between the total unsigned magnetic flux and the total radiance in the field of view. The residuals between the simple linear model and the observed total radiance is also shown.

we illustrate the robustness of this assumption for EUV emission from small regions of the solar disk so we can discuss this in terms of a flux-radiance relationship. We also show that it is possible to use the magnetic field integrated over the disk as a replacement for other proxies of solar activity.

The relationship between the radiance and the unsigned magnetic flux is illustrated in Fig. 2. The upper panels in this figure show the central portion of the solar disk from several MDI magnetograms (top) and EIT images at 304 Å (bottom) from October 7 through November 5, 2003. This period is one of particularly high solar activity and is a challenge for any modeling. The two panels on the lower left show the times series of the EIT 304 Å emission and unsigned magnetic flux integrated over this region ( $\Phi_r = \sum |B_r| \cdot A$ , where  $B_r$  is the flux density in Gauss in a pixel and  $A$  is the pixel area). The strong correlation is clearly shown by these two time series. A scatter plot of these values is shown in the lower right portion of Fig. 2 along with

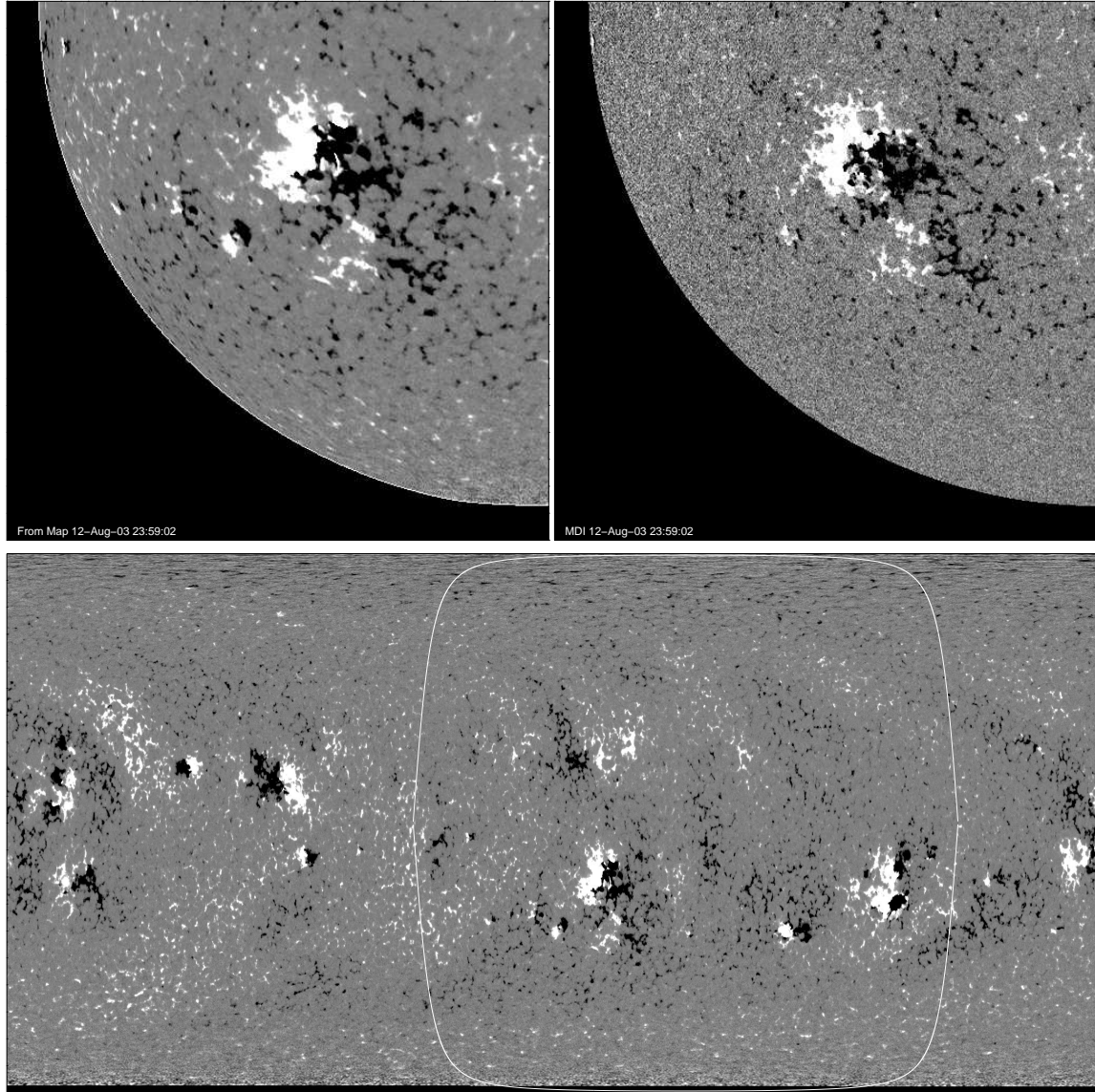


Figure 3: A comparison of a radial magnetogram derived from a Carrington map (upper left) and a standard, line-of-sight, synoptic MDI magnetogram (upper right). The use of the Carrington map (bottom panel) greatly improves the signal-to-noise at the limb but at the expense of using information that is not necessarily current.

a simple linear fit. As indicated, the correlation is very high ( $r = 0.96$ ). Finally, the residuals between the observations and linear fit are plotted and give an estimate of the uncertainty in the fit ( $\sigma = 6.8\%$ ).

One limitation in using magnetograms as the basis for irradiance modeling is that magnetic fields in the photosphere are predominately radial and are difficult to measure at the limb. This is illustrated in Fig. 3, where we show a quadrant from an instantaneous MDI magnetogram. To address this issues we make use of Carrington synoptic maps, two dimensional  $(\theta, \phi)$  maps of the solar surface that are built up as features rotate past central meridian, to project information accumulated at disk center to the limb. Since magnetic fields generally evolve slowly this usually does not introduce significant error. An example Carrington map and corresponding magnetogram is also shown in Fig. 3.

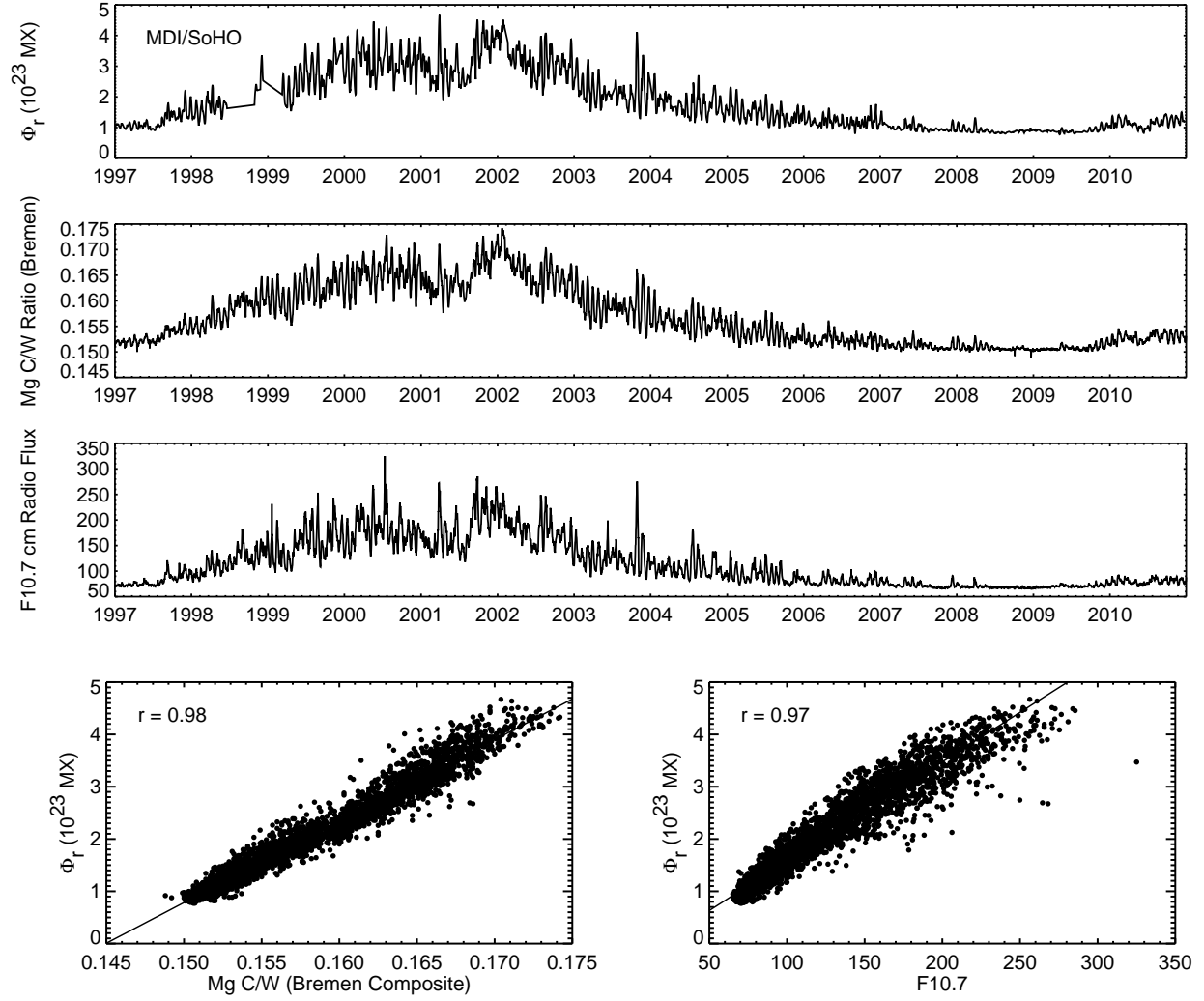


Figure 4: Three proxies for solar activity  $\Phi_r$  (the disk-integrated  $|B_r|$  derived from MDI observations), the Mg core-to-wing ratio, and the 10.7 cm radio flux. The bottom panels show the regression between the disk-integrated Br and both 10.7 and Mg core-to-wing.

We have computed daily sums of  $\Phi_r$  integrated over the disk using synoptic maps from both MDI and the Helioseismic and Magnetic Imager (HMI) on the Solar Dynamics Observatory (SDO). The time series for the integrated  $\Phi_r$  are shown in Fig. 4 and 5 along with F10.7 and the Mg core-to-wing ratio. We also show the correlation between these proxies and  $\Phi_r$ . The correlation between  $\Phi_r$  and existing proxies for solar activity is high ( $r = 0.98$  for the Mg core-to-wing ratio and  $r = 0.97$  for F10.7), suggesting that  $\Phi_r$  could be used as replacement for these other proxies for solar activity.

The use of  $\Phi_r$  has a number of advantages over existing proxies. Synoptic magnetic field measurements are made by a number of instruments located either in space or at ground based observatories, thus providing redundant sources of information. Magnetic field measurements derived from Zeeman splitting depend on the differences in polarization at the same wavelength and are much less susceptible to long-term instrumental degradation. Since this proxy uses spatially resolved observations it can be tailored to the center-to-limb behavior of the emission of interest. Finally, the evolution of solar magnetic fields is generally well understood, making a proxy derived from measured magnetic fields potentially much easier

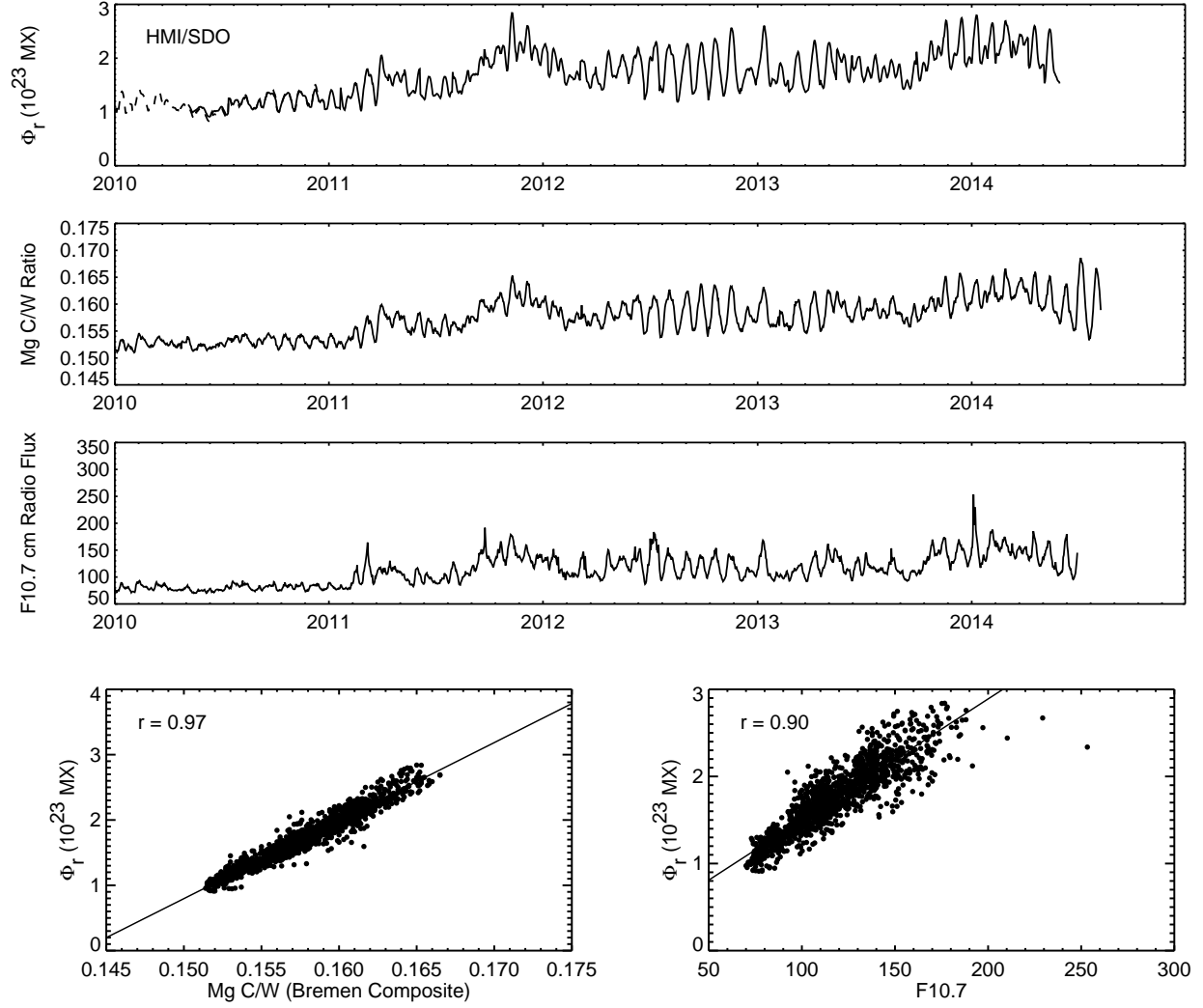


Figure 5: The same as Fig. 4, but for  $\Phi_r$  derived from HMI observations of the magnetic field. The dotted line in the top panel shows the results from MDI for this period.

to forecast.

#### 4. Magnetic Flux Transport and Forecasting Magnetic Activity

The evolution of magnetic fields on the solar surface can be described by a two-dimensional convection-diffusion equation:

$$\frac{\partial B_r}{\partial t} = -\omega(\theta) \frac{\partial B_r}{\partial \phi} - \frac{1}{R_\odot \sin \theta} \frac{\partial}{\partial \theta} [v(\theta) B_r \sin \theta] + \kappa \nabla_\perp^2 B_r + S \quad (2)$$

where  $\omega(\theta)$  is the differential rotation rate,  $v(\theta)$  is the meridional flow velocity,  $\kappa$  is the diffusion rate, and  $S$  is the flux emergence rate [see 10, for a review].

We have implemented a discrete flux transport model similar to those of *Wang and Sheeley* [17], *Worden and Harvey* [19], and *Schrijver* [13], but that is optimized for use with new observations at higher spatial



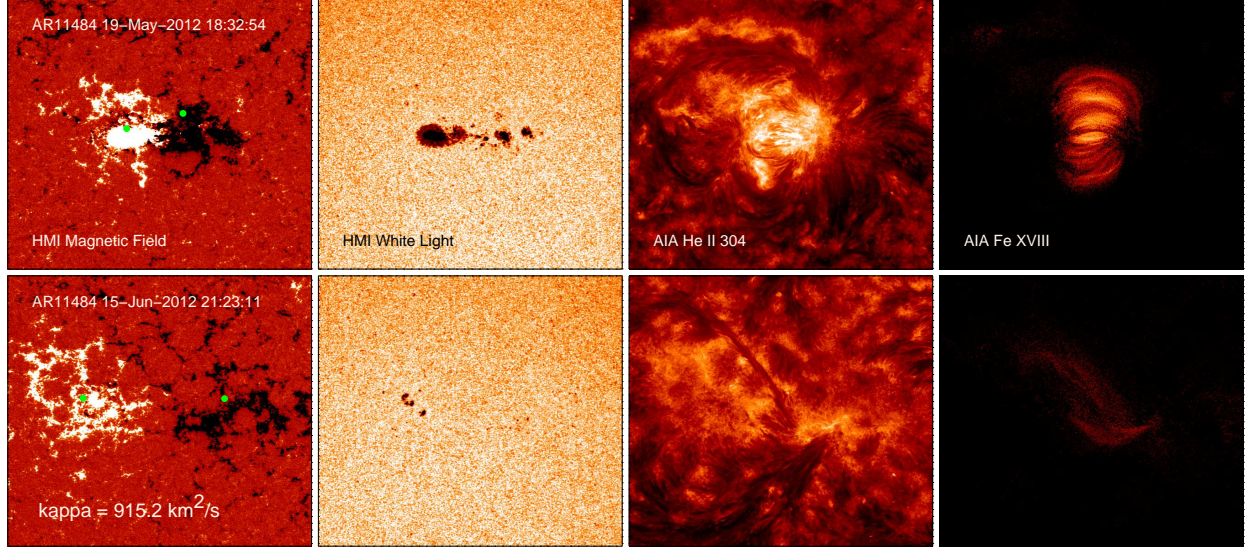


Figure 6: An example measurement of the diffusion coefficient  $\kappa$  using SDO observations of an active region observed over a solar rotation. The green dots plotted on the magnetograms (far left) indicate the centroids for each polarity and show the increase in the separation between the polarities over this time.

resolution. We are in the process of validating this model against previous model results. We use the classic rates for differential rotation

$$\omega(\theta) = 13.38 - 2.30 \sin^2 \theta - 1.62 \sin^4 \theta \quad \text{deg day}^{-1} \quad (3)$$

and meridional flow

$$v(\theta) = 12.9 \sin 2\theta + 1.4 \sin 4\theta \quad \text{m s}^{-1}. \quad (4)$$

At each time step each element of flux undergoes advection and diffusion. Flux cancellation occurs when elements of opposite polarity move into the same grid cell.

Historically there have been large differences in the diffusion rate ( $\kappa$ ) inferred from observations, with values ranging from 200 to 1000  $\text{km}^2 \text{s}^{-1}$ . *Schrijver* [13] speculated that this results from a dependence of the diffusion rate on the magnitude of the flux, perhaps resulting from a magnetoconvective coupling between the surface and lower layers of the solar atmosphere. We are making a new measurement of the supergranular diffusion coefficient from SDO observations. The nearly continuous observations from SDO allow for a systematic study of active region evolution using hundreds of examples. One such measurement is illustrated in Fig. 6 where we show the magnetic field, the white light photosphere, the upper transition region, and the high temperature corona. Our initial measurements support relatively large values for the diffusion coefficient, but more analysis is needed.

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